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than one second of arc. After passing the Sun, the planet was not looked for by daylight until August 22d, at 10 A.M., when it was, of course, very easily seen. R. G. AITKEN.

DISCOVERY OF THREE HUNDRED DOUBLE STARS.

During the years 1898-99, I was engaged in remeasuring all the double stars discovered at Pulkowa, principally from 1841 to 1843, by OTTO STRUVE. In the course of this work I discovered new companions to a few of these stars. I also frequently measured other pairs which happened to be in the vicinity of those found by OTTO STRUVE, and some of these miscellaneous pairs have proved to be new.

When the work on the Pulkowa stars drew towards a close, so that considerable gaps existed in my programme, I began to examine large numbers of stars in various parts of the sky, particularly towards the south, in the hope of discovering new pairs. This work has now been carried on systematically for about fourteen months, resulting in the discovery of three hundred new double stars, having distances under five seconds of arc.

Sets of measures of these stars have been obtained as speedily as circumstances have permitted. Observations of one hundred pairs were printed in the *Astronomical Journal*, No. 480, and observations of a second hundred have been forwarded to the same periodical for publication. Sets of measures of the third hundred have not yet been completed, though most of them have been observed on at least one night.

With respect to the distances between their components, these new pairs have the following classification:

0".25 or less,	18 pairs.
0.26 to 0".50,	41 "
0.51 to 1.00,	63 "
1.01 to 2.00,	80 "
2.01 to 5.00,	98 "
	<hr/> 300

Many of these stars are difficult to measure, either on account of the closeness of their components or the faintness of their companions. On this account I have generally observed them with the 36-inch telescope. Some of the close pairs, even some of those whose distances do not exceed 0".25, have been discov-

ered with the 12-inch telescope. But measures of such pairs have not been made with this instrument.

It would have been an easy matter to have increased the number of discoveries considerably by including pairs having distances somewhat greater than 5". This, however, has not seemed desirable, and only a few such pairs have been measured, and the results for one only published. W. J. HUSSEY.

September 18, 1900.

ELLIPTIC ELEMENTS OF COMET 1896 V, GIACOBINI.

In November, 1896, I computed elements of the orbit of this comet from places corresponding to the dates 1896, September 8th, October 5th and 29th. The elements then obtained did not accurately represent the observed path of the comet during the latter part of its apparition. About a year ago I found the origin of the discrepancy in a slightly erroneous position of the comparison-star used for the last observation, and at that time computed the system of elements given below. These elements are based on normal places having the dates 1896, September 8th, October 29th, and 1897, January 4th. The interval included between the extreme dates is only four days less than the apparition period of the comet.

Epoch 1896 November 1.5 G. M. T.

$$M = 0^{\circ} \ 39' \ 56''.5$$

$$\Omega = 193 \ 28 \ 47.4$$

$$i = 11 \ 21 \ 14.7 \left. \vphantom{\begin{matrix} \Omega \\ i \end{matrix}} \right\} 1896.0$$

$$\omega = 140 \ 33 \ 43.4$$

$$\phi = 35 \ 42 \ 7.6$$

$$\log e = 9.766094$$

$$\log a = 0.542898$$

$$\mu = 544''.0765$$

$$\text{Period} = 6.52 \text{ years.}$$

W. J. HUSSEY.

PROPER MOTION OF $\text{O}\Sigma \ 165, 45 \text{ GEMINORUM.}$

OTTO STRUVE discovered a faint companion of *45 Geminorum* about 1842, and first measured its position in 1847. Misled by erroneous observations made in 1877, he thought that the change in the relative positions of the two stars was due to orbital motion. This, however, is not the case. It results entirely from the proper motion of the principal star. This is clearly estab-